

Holographic superconductors from gauged supergravity

Francesco Aprile,^a Diederik Roest^b and Jorge G. Russo^{a,c}

^a*Institute of Cosmos Sciences and Estructura i Constituents de la Materia,
Facultat de Física, Universitat de Barcelona,
Av. Diagonal 647, 08028 Barcelona, Spain*

^b*Centre for Theoretical Physics, University of Groningen,
Nijenborgh 4, 9747 AG Groningen, The Netherlands*

^c*Institució Catalana de Recerca i Estudis Avançats (ICREA),
Pg. Lluís Companys, 23, 08010 Barcelona, Spain*

E-mail: aprile@ecm.ub.es, d.roest@rug.nl, jrusso@ub.edu

ABSTRACT: We consider minimal setups arising from different truncations of $\mathcal{N} = 8$ five-dimensional $SO(6)$ gauged supergravity to study phase transitions involving spontaneous breaking of any of the $U(1)$ symmetries in $U(1) \times U(1) \times U(1) \subset SO(6)$. These truncations only keep the three relevant vector fields, four complex scalar fields carrying $U(1)$ charges, plus two neutral scalar fields required by consistency. By considering thermal ensembles with different fixed $U(1)$ charge densities and solving the complete equations including the full back-reaction, in some cases we find instabilities towards the formation of hairy black holes, which lead to second order transitions, resulting from a thermodynamical competition between different sectors. We argue that these should be the dominant thermodynamical instabilities in the full ten-dimensional type IIB theory. In other cases we find unstable branches of hairy black holes that extend to temperatures above a critical temperature (“retrograde condensation”). The results can be used as a first step to understand new aspects of the phase diagram of large N $\mathcal{N} = 4$ $SU(N)$ super Yang-Mills theory with fixed charge densities.

KEYWORDS: Gauge-gravity correspondence, Black Holes in String Theory, Supergravity Models, Holography and condensed matter physics (AdS/CMT)

ARXIV EPRINT: [1104.4473](https://arxiv.org/abs/1104.4473)

Contents

1	Introduction	1
2	Simple truncations of $D = 5$ maximal gauged supergravity	4
2.1	A supersymmetric truncation	4
2.2	Non-supersymmetric truncation I: keeping scalars in the $\mathbf{20}_S$	6
2.3	Non-supersymmetric truncation II: keeping scalars in the $\mathbf{20}_A$	9
3	Thermodynamics and condensed matter applications	12
3.1	Equations of motion and asymptotic behavior	12
3.2	The uncondensed phase	15
3.3	The condensed phase	18
4	Condensation from other sectors	27
5	Concluding remarks	28

1 Introduction

One of the important problems in holographic superconductivity [1, 2]¹ is to disclose the precise dictionary between gravity and the condensed matter system. This requires what is called a “top-down” approach, starting with ten-dimensional string theory or M-theory and a brane construction so that the field theory undergoing the phase transition can be explicitly understood [5–9]. These theories contain many degrees of freedom whose dual operator may condense and break a $U(1)$ symmetry, thus leading to superconductivity. Identifying the most relevant mode, the one which is dual to the operator that condenses first and dominates the thermodynamics, is a priori a complicated task. However, an important clue lies in the empirical observation that operators with lower dimensions and higher R-charges generically have a higher critical temperature, and hence are prone to condense earlier as the temperature is lowered (see e.g. [5, 6, 10]). Therefore, a first step in the top-down approach would be the construction of consistent truncations that include the lightest modes in the mass spectra arising from the higher dimensional theory. Here we will consider the spectra of theories with maximal supersymmetry that are relevant for holography, where the lightest modes are scalar fields, which can have negative masses without inducing perturbative instabilities in AdS.

The three maximal supergravities with Anti-de Sitter vacua, in $D = 7, 5$ and 4 , share a number of features. First of all, they all come from Freund-Rubin reductions of ten- or eleven-dimensional parent theories over an S^4 , S^5 and S^7 , respectively. As a consequence

¹For reviews and a more complete list of references, see [3, 4].

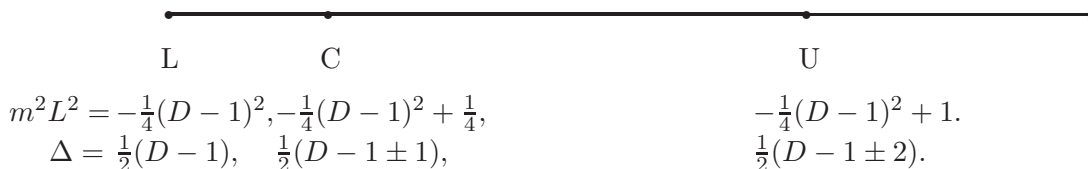


Figure 1. The different values of allowed (negative) masses and dimensions in Anti-de Sitter. Masses on the bold line allow for two physically relevant dimensions.

of the S^{n-1} reduction, the gauge group is always $SO(n)$. This is embedded in an $SL(n)$ subgroup of the full global symmetry group, which is $SL(5)$, E_6 and E_7 , respectively. In view of the above, it is important to elucidate the mass spectra of the scalar fields of these theories.

In the case of AdS backgrounds, a lower bound on scalar masses is set by the Breitenlohner-Freedman (BF) bound at [11]

$$m_L^2 L^2 = -\frac{1}{4}(D-1)^2, \tag{1.1}$$

where the scale L of AdS is set by the scalar potential via the relation $VL^2 = -(D-1)(D-2)$, with the scalar potential evaluated in a critical point. In addition to the above, there are two other interesting values of the masses slightly above the BF bound [12]. The first is at

$$m_C^2 L^2 = -\frac{1}{4}(D-1)^2 + \frac{1}{4}, \tag{1.2}$$

and corresponds to a conformally coupled scalar field. Finally, there is

$$m_U^2 L^2 = -\frac{1}{4}(D-1)^2 + 1. \tag{1.3}$$

This is the upper bound for scalar masses that still allow for two different boundary conditions. The dimension Δ of the dual operator is related by $m^2 L^2 = \Delta(\Delta - D + 1)$. The unitarity bound implies that the lowest possible dimension is $\Delta = \frac{1}{2}(D-3)$, corresponding to the upper bound on the mass. Upon increasing the dimension, the corresponding mass first goes down to the Breitenlohner-Freedman lower bound, after which it goes up to plus infinity. Both the masses and dimensions are illustrated in figure 1.

All three theories have an Anti-de Sitter vacuum in the origin, which is maximally supersymmetric and preserves the $SO(n)$ gauge group. On account of this, the scalars are organised in $SO(n)$ irreps with particular masses. It turns out that the three theories under consideration exactly have masses lying at the three special mass values listed above [12]. In particular, the seven-dimensional theory has a 14-dimensional irrep of scalars with

$$D = 7 : \quad m^2 L^2 = -8, \tag{1.4}$$

corresponding to the upper bound. In contrast, the four-dimensional theory has two 35-dimensional irreps of scalars with

$$D = 4 : \quad m^2 L^2 = -2, \tag{1.5}$$

corresponding to a conformally coupled scalar field. Finally, the five-dimensional theory has

$$D = 5 : \quad m^2 L^2 = -4 \quad 20 \text{ scalars} , \quad (1.6)$$

$$m^2 L^2 = -3 \quad 20 \text{ scalars} , \quad (1.7)$$

$$m^2 L^2 = 0 \quad 2 \text{ scalars} . \quad (1.8)$$

Thus 20 scalars saturate the BF bound, whereas another set of 20 scalars saturate the unitarity bound (more details will follow in the next section). Therefore, the two maximal supergravities coming from M-theory have masses that satisfy the Breitenlohner-Freedman bound, while the $D = 5$ theory coming from IIB has a number of scalars that saturate the bound.

The connection between type IIB string theory on $\text{AdS}_5 \times \text{S}^5$ to large- N $\mathcal{N} = 4$ $\text{SU}(N)$ super Yang-Mills theory is the best understood AdS/CFT duality. In view of this, we will focus on the five-dimensional supergravity case. The relevant five-dimensional description is given by $\mathcal{N} = 8$ $\text{SO}(6)$ gauged supergravity [13, 14] and truncations thereof, which will be here used as a basic setting to understand superconductivity (or, more precisely, superfluidity) in $\mathcal{N} = 4$ super Yang-Mills theory. This theory has $\text{SO}(6)$ global R-symmetry, which contains $\text{U}(1) \times \text{U}(1) \times \text{U}(1)$ as its maximal Abelian subgroup. In this context, one can thus consider a canonical ensemble with fixed charge densities $\hat{\rho}_1, \hat{\rho}_2, \hat{\rho}_3$ associated with each $\text{U}(1)$ group. The problem is then to understand what is the phase diagram of the theory as the temperature is gradually lowered from high values. On the gravity side, the high temperature thermodynamics is dominated by the STU black hole with charges proportional to $\hat{\rho}_1, \hat{\rho}_2, \hat{\rho}_3$. As the temperature gets lower than some critical value, some scalar operators are expected to condense, possibly breaking $\text{U}(1)$ symmetries. Understanding this condensation process in the full type IIB superstring theory context is obviously complicated. However, the operators which should condense first are those with low dimensions, and hence are dual to light modes of $D = 5$ $\mathcal{N} = 8$ $\text{SO}(6)$ gauged supergravity. We will investigate a number of such modes for various configurations of black hole charges.

This paper is organised as follows. In the next section we will discuss maximal supergravity in five dimensions, together with a truncation to minimal supergravity and two non-supersymmetric truncations (sectors I and II). All of these theories have a $\text{U}(1)^3$ gauge group and a number of scalar fields. In section 3 we will consider charged black hole solutions of this theory. In particular, we will first review the black hole solutions corresponding to the uncondensed phase. Subsequently the possibility of black holes with scalar hair, corresponding to a condensed phase, is investigated in the two non-supersymmetric sectors I and II. We will show which phase is thermodynamically dominant as a function of temperature, and relate our results to previous literature. In section 4 we discuss possible condensation from other sectors not included in the previous truncations. Finally, some concluding remarks are given in section 5.

Note added. We have been informed that a complementary study of black hole solutions with hair of $D = 4$ gauged supergravity will appear in [15], in concert with this paper. A study of models related to $\mathcal{N} = 4$ SYM with chemical potentials from the field theory side will appear in [16].

2 Simple truncations of $D = 5$ maximal gauged supergravity

2.1 A supersymmetric truncation

Maximal supergravity in $D = 5$ has a global E_6 symmetry [17]. In particular, the scalar fields span the coset $E_6/\text{Usp}(8)$. However, the presence of an $\text{SO}(6)$ gauging breaks this symmetry. Instead, it is convenient to arrange the fields in irreps of the $\text{SL}(2) \times \text{SL}(6)$ maximal subgroup or its compact subgroup.

In terms of the $\text{SL}(2) \times \text{SL}(6)$ maximal subgroup, the isometries of E_6 decompose as

$$\text{scalar isometries: } (\mathbf{1}, \mathbf{35}) \oplus (\mathbf{2}, \mathbf{20}) \oplus (\mathbf{3}, \mathbf{1}). \tag{2.1}$$

The first and last term correspond to $\text{SL}(6)$ and $\text{SL}(2)$, respectively, while the additional irrep in the middle combines these into E_6 . Physical scalars correspond to the non-compact isometries, of which there are (of course) 42. In terms of $\text{SO}(2) \times \text{SO}(6)$, these are given by

$$\text{scalars: } (\mathbf{1}, \mathbf{20}_S) \oplus (\mathbf{2}, \mathbf{20}_A)^+ \oplus (\mathbf{2}, \mathbf{1}), \tag{2.2}$$

where the $\mathbf{20}_S$ denotes a symmetric rank-2 and the $\mathbf{20}_A$ is an anti-symmetric rank-3 tensor. Moreover, the $+$ corresponds to an imaginary self-dual condition.² Surprisingly, the supermultiplet structure in $D = 5$ is such that not all scalars have identical masses. Instead, the three irreps listed above have

$$m^2 L^2 = -4, \quad -3, \quad 0, \tag{2.3}$$

respectively [18]. The first two of these exactly correspond to the lower and upper bounds defined in the introduction.

The vector bosons are in the anti-symmetric representation of $\text{SL}(6)$:

$$\text{vectors: } (\mathbf{1}, \mathbf{15}), \tag{2.4}$$

which are of course massless. The theory also contains twelve two-forms:

$$\text{two-forms: } (\mathbf{2}, \mathbf{6}). \tag{2.5}$$

In the ungauged theory, the two-forms can be dualised to vectors and combine to form the $\mathbf{27}$ irrep of E_6 ; however, due to the $\text{SO}(6)$ gauging this is no longer possible in the gauged supergravity. Instead, the two-forms acquire a mass term and have $m^2 L^2 = +1$ [18].

As for the fermions, these consists of 4 gravitini and 24 dilatini. The former are in the spinorial representation of $\text{SO}(6)$ and have masses $m^2 L^2 = +\frac{3}{2}$. The latter are in $\mathbf{4}$ and $\mathbf{20}$ representations with masses $-\frac{3}{2}$ and $-\frac{1}{2}$, respectively.

²In more detail, the $(\mathbf{2}, \mathbf{20})$ transforms as a doublet of anti-symmetric three-forms of $\text{SL}(2) \times \text{SO}(6)$. Branching this to $\text{SO}(2) \times \text{SO}(6)$ one obtains two irreps: the (anti-)imaginary self-dual combinations $(\mathbf{2}, \mathbf{20}_A)^\pm$. One of these correspond to the compact generators of E_6 , and hence is modded out, while the other combination is retained and corresponds to physical scalars.

We now consider a consistent truncation of the maximal SO(6) supergravity by modding out by a $\mathbb{Z}_2 \times \mathbb{Z}_2$ symmetry with generators³

$$\begin{pmatrix} -\mathbb{I}_4 & \\ & +\mathbb{I}_2 \end{pmatrix}, \quad \begin{pmatrix} +\mathbb{I}_2 & \\ & -\mathbb{I}_4 \end{pmatrix}, \tag{2.6}$$

acting on the fundamental representation of SL(6). In fact, this corresponds to a supersymmetric truncation, as can be seen from the following. The SO(6) transformations translate into the following SU(4) transformation:

$$\begin{pmatrix} +1 & & & \\ & -1 & & \\ & & +1 & \\ & & & -1 \end{pmatrix}, \quad \begin{pmatrix} +1 & & & \\ & -1 & & \\ & & -1 & \\ & & & +1 \end{pmatrix}, \tag{2.7}$$

where we have related the fundamental of SO(6) to the anti-symmetric representation of SU(4) via invariant tensors, which are a permutation of the 't Hooft symbols. From (2.7) it follows that this truncation preserves one out of the four original supersymmetries (thus there are eight preserved supercharges).

The resulting field content is as follows. Firstly, all two-forms are removed by the truncation. Secondly, only three vector bosons survive, associated with the maximal Abelian subgroup $U(1) \times U(1) \times U(1)$ in SO(6). Finally, out of the three irreps of scalar isometries, we retain respectively eleven, sixteen and three isometries. The corresponding numbers of physical scalars are eight, eight and two. Similarly, on the fermionic side we find a single gravitino and six spin-1/2 fields. These fields get organized in the following multiplets of minimal supersymmetry in $D = 5$:

- Gravity: the graviton, a gravitino and a vector,
- Vector: a vector, a gaugino and a real scalar,
- Hyper: a hyperino and four real scalars.

Therefore the resulting $\mathcal{N} = 1$ theory contains, in addition to the gravity multiplet, two vector and four hyper multiplets. The scalar manifold is given by

$$SO(1,1)^2 \times \frac{SO(4,4)}{SO(4) \times SO(4)}, \tag{2.8}$$

where the first factor is the very special Kahler geometry spanned by the vector multiplet, and the second factor is the Quaternionic-Kahler manifold of the hyper sector. The same result was found in [20].

In what follows we will consider two different non-supersymmetric truncations, that share the two dilatons of the vector multiplets but pick out two completely orthogonal $(SL(2)/SO(2))^4$ scalar submanifolds of the hyper sector. In both truncations we will pick out the scalar mode(s) that are relevant to describe hair in various black hole backgrounds.

³This truncation was studied in [19] and more recently in [20].

2.2 Non-supersymmetric truncation I: keeping scalars in the 20_S

We now make the first subsequent truncation, to bring the theory to a more manageable form. In addition to the $\mathbb{Z}_2 \times \mathbb{Z}_2$ generators acting on the fundamental of $SL(6)$, we also mod out by the $SO(2)$ generator $-\mathbb{I}_2$ acting on the fundamental of $SL(2)$. The latter has the following two implications:

- Supersymmetry is now lost: the corresponding transformation in terms of $U(4)$ leaves none of the four supersymmetries invariant.
- The $(\mathbf{2}, \mathbf{20}_A)$ irrep is removed by the truncation and one is only left with the $SL(2) \times SL(6)$ scalar manifold.

This corresponds to the five-sphere reduction of IIB supergravity with only the ten-dimensional metric and five-form retained.⁴ Due to the second point, the scalar potential reduces to a universal formula, that is also valid for the $SL(n)/SO(n)$ subgroup of the other AdS maximal supergravities. For the vectors and the $SL(n)$ subgroup of scalars, which we will parametrise by a symmetric matrix T , the Lagrangian reads⁵ [22, 23]

$$\mathcal{L} = \sqrt{-g} \left[R - \frac{1}{4} \text{Tr}[T^{-1} D_\mu T T^{-1} D^\mu T] - \frac{1}{4} \text{Tr}[T^{-1} F_{\mu\nu} T^{-1} F^{\mu\nu}] - V \right], \quad (2.9)$$

where the covariant derivatives on the scalars read

$$D_\mu T = \partial_\mu T + g_0 (A_\mu T - T A_\mu). \quad (2.10)$$

Finally, the scalar potential is given by

$$V = \frac{1}{2} g_0^2 (2\text{Tr}[TT] - \text{Tr}[T]^2). \quad (2.11)$$

The coupling constant g_0 sets the scale for the AdS radius L . In $D = 5$ this relation reads $g_0 L = 1$. An important feature of this truncation in five dimensions is that it retains the $SL(6)/SO(n)$ scalar fields in the scalar potential with masses at the Breitenlohner-Freedman bound, which have a high chance to condense due to the low dimension of their dual operators. Finally, note that the additional $SL(2)/SO(2)$ scalars do not appear in these expressions and are massless and neutral.

The combined action of the $SO(6)$ transformations (2.6) with the $SO(2)$ element $-\mathbb{I}_2$ leads to a field content containing three vectors and ten scalars. The latter span the non-compact part of the remaining symmetry group

$$SO(1,1)^2 \times SL(2)^4. \quad (2.12)$$

The part of the $SL(6)$ scalar manifold that is invariant under the $\mathbb{Z}_2 \times \mathbb{Z}_2$ symmetries (2.6) can be parametrised by

$$T_{mn} = \begin{pmatrix} X_1 \mathcal{M}_1 & & \\ & X_2 \mathcal{M}_2 & \\ & & X_3 \mathcal{M}_3 \end{pmatrix}, \quad (2.13)$$

⁴This truncation was previously studied in [21].

⁵In addition there can be topological Wess-Zumino terms for the vectors fields. We have not included these as these will not play any role in what follows.

where

$$\mathcal{M}_i = \begin{pmatrix} \cosh \eta_i + \sinh \eta_i \cos \theta_i & \sinh \eta_i \sin \theta_i \\ \sinh \eta_i \sin \theta_i & \cosh \eta_i - \sinh \eta_i \cos \theta_i \end{pmatrix}, \quad (2.14)$$

and $X_1 X_2 X_3 = 1$. A parametrisation of the latter is

$$X_1 = e^{\varphi_1 - \varphi_2}, \quad X_2 = e^{\varphi_1 + \varphi_2}, \quad X_3 = e^{-2\varphi_1}. \quad (2.15)$$

We will use the same form (2.14) for the separate $SL(2)$ scalar, parametrised by \mathcal{M}_4 . The four scalars are collected in $\eta_a = (\eta_i, \eta_4)$, and similar for θ_a .

The Lagrangian of the remaining bosonic fields is then given by

$$\mathcal{L} = \sqrt{-g} \left(R - \sum_{i=1}^3 X_i^{-2} \left[\frac{1}{2} (\partial X_i)^2 + \frac{1}{4} F^{i,\mu\nu} F_{\mu\nu}^i \right] - \frac{1}{2} \sum_{a=1}^4 [(\partial \eta_a)^2 + \sinh^2 \eta_a (\partial \theta_a + L^{-1} q_{ai} A_i)^2] - V \right), \quad (2.16)$$

where we used the notation $A_i = (A^{12}, A^{34}, A^{56})$. These span the $U(1)^3$ remaining part of the gauge group. The scalar charges with respect to it are given by

$$q_{1i} = (2, 0, 0), \quad q_{2i} = (0, 2, 0), \quad q_{3i} = (0, 0, 2), \quad q_{4i} = (0, 0, 0). \quad (2.17)$$

Finally, despite the absence of supersymmetry in this truncation, the resulting scalar potential can be written in terms of a superpotential

$$V = \frac{1}{2L^2} \left[\sum_{a=1}^4 \left(\frac{\partial W}{\partial \eta_a} \right)^2 + \sum_{i=1}^3 X_i^2 \left(\frac{\partial W}{\partial X_i} \right)^2 \right] - \frac{1}{3L^2} W^2, \quad (2.18)$$

where

$$W = \sum_{a=1}^4 q_{ai} X_i \cosh(\eta_a). \quad (2.19)$$

Note that the first three angular scalars θ_i are pure gauge and can be set equal to zero by a $U(1)^3$ transformation. It can be checked that the origin of moduli space is an extremum. Similarly, the eigenvalues of the Hessian of the scalar potential are given by

$$m_a^2 L^2 = - \sum_{i=1}^3 q_{ai} q_{ai}. \quad (2.20)$$

These are equal for the first three scalars η_i and given by -4 . This indeed corresponds to all scalar masses saturating the Breitenlohner-Freedman bound. Finally, the last scalar η_4 drops out of the scalar potential on account of being neutral.

In what follows, three further truncations of this model will be considered. In particular, we will consider subsectors where a number of gauge vectors are identified, and focus on the dynamics of at most two scalar fields. As these truncations will later be used to study the emergence of hair in charged black holes, the number of gauge fields will determine the number of black hole charges. These will be used to label the various cases.

One single charge (I). First of all, we consider a truncation of the three-block model to a sector with just a single BH charge. In this case we augment the $\mathbb{Z}_2 \times \mathbb{Z}_2$ of (2.6) to a full $\text{SO}(4)$ that is orthogonal to the diagonal $\text{SO}(2)$. In other words, we will now consider a truncation of maximal supergravity, based on the decomposition of $\text{SO}(6)$ into $\text{SO}(4) \times \text{SO}(2)$. In terms of special unitary groups this reads $\text{SU}(4)$ into $\text{SU}(2) \times \text{SU}(2) \times \text{U}(1)$, where the two $\text{SU}(2)$ factors are block-diagonal generalisations of (2.7). The different irreps of $\text{SO}(6) \simeq \text{SU}(4)$ then split up as

$$\begin{aligned}
 \mathbf{4} &\rightarrow \mathbf{2}_{+1/2} \oplus \mathbf{2}_{-1/2}, \\
 \mathbf{6} &\rightarrow \mathbf{4}_0 \oplus \mathbf{1}_{+1} \oplus \mathbf{1}_{-1}, \\
 \mathbf{15} &\rightarrow \mathbf{6}_0 \oplus \mathbf{4}_{+1} \oplus \mathbf{4}_{-1} \oplus \mathbf{1}_0, \\
 \mathbf{20}_S &\rightarrow \mathbf{9}_0 \oplus \mathbf{4}_{+1} \oplus \mathbf{4}_{-1} \oplus \mathbf{1}_{+2} \oplus \mathbf{1}_0 \oplus \mathbf{1}_{-2},
 \end{aligned}
 \tag{2.21}$$

in terms of $\text{SO}(4)$ irreps with an $\text{SO}(2)$ weight. Subsequently we only retain the singlets of $\text{SO}(4)$. This amounts to one vector, corresponding to the remaining $\text{SO}(2) \sim \text{U}(1)$ gauge group, and three scalar fields. The latter split up in one real scalar, and one complex scalar field with $\text{U}(1)$ charge $q = +2$. Furthermore, from the decomposition of the $\mathbf{4}$ one can see that this truncation does not preserve any supersymmetry.

Instead of the diagonal $\text{SO}(2)$, we will use a different but equivalent embedding, where the $\text{SO}(2)$ acts on the last two indices. Consequently, the $\text{SO}(4)$ acts on the first four indices. In terms of scalars, this truncation amounts to the Ansatz

$$\eta_a = (0, 0, \eta, 0), \quad \theta_a = (0, 0, \theta, 0), \quad \varphi_2 = 0.
 \tag{2.22}$$

Similarly, for the vector fields this requires $A_i = (0, 0, A)$. The complete Lagrangian takes the form

$$\mathcal{L} = \sqrt{-g} \left[R - 3(\partial\varphi_1)^2 - \frac{1}{2}(\partial\eta)^2 - \frac{1}{4}e^{4\varphi_1} F_{\mu\nu} F^{\mu\nu} - \frac{1}{2} \sinh^2 \eta (\partial\theta + 2L^{-1}A)^2 - V \right],
 \tag{2.23}$$

with

$$V = \frac{1}{L^2} e^{4\varphi_1} (-4 - 2e^{-12\varphi_1} + 2e^{-12\varphi_1} \cosh^2 \eta - 8e^{-6\varphi_1} \cosh \eta).
 \tag{2.24}$$

Two equal charges (I). Similarly, one can consider a truncation to a subsector that can carry two of the three BH charges. In order to truncate to a smaller number of scalars, one can mod out by the $\text{SO}(6)$ element

$$\begin{pmatrix} & +\mathbb{I}_2 & \\ -\mathbb{I}_2 & & \\ & & +1 \\ & & & -1 \end{pmatrix}.
 \tag{2.25}$$

This has the following effect on the remaining scalars:

$$\eta_a = (\eta, \eta, 0, 0)/\sqrt{2}, \quad \theta_a = (\theta, \theta, 0, 0), \quad \varphi_2 = 0.
 \tag{2.26}$$

Furthermore, consistent with the above truncation, we take the gauge vectors equal to

$$A_i = (A, A, 0)/\sqrt{2}. \tag{2.27}$$

This implies that the Lagrangian reduces to

$$\mathcal{L} = \sqrt{-g} \left[R - 3(\partial\varphi_1)^2 - \frac{1}{2}(\partial\eta)^2 - \frac{1}{4}e^{-2\varphi_1}F_{\mu\nu}F^{\mu\nu} - \sinh^2(\eta/\sqrt{2})(\partial\theta + \sqrt{2}L^{-1}A)^2 - V \right], \tag{2.28}$$

with the scalar potential given by

$$V = -\frac{4}{L^2}e^{-\varphi_1}(2 \cosh(\eta/\sqrt{2}) + e^{3\varphi_1}). \tag{2.29}$$

Three equal charges (I). Finally, a very simple sector is obtained by further restricting to the ‘isotropic’ truncation, where the three $SL(2)$ copies are identified with each other. Group-theoretically this can be obtained by modding out by $SO(6)$ elements

$$\left(\begin{array}{ccc} & +\mathbb{I}_2 & \\ -\mathbb{I}_2 & & \\ & & \mathbb{I}_2 \end{array} \right), \quad \left(\begin{array}{ccc} & +\mathbb{I}_2 & \\ & & +\mathbb{I}_2 \\ & -\mathbb{I}_2 & \end{array} \right). \tag{2.30}$$

This requires the further restrictions on the scalar fields

$$\eta_a = (\eta, \eta, \eta, 0)/\sqrt{3}, \quad \theta_a = (\theta, \theta, \theta, 0), \quad \varphi_1 = \varphi_2 = 0. \tag{2.31}$$

Similarly we set

$$A_i = (A, A, A)/\sqrt{3}. \tag{2.32}$$

The resulting Lagrangian is

$$\mathcal{L} = \sqrt{-g} \left[R - \frac{1}{4}F^{\mu\nu}F_{\mu\nu} - \frac{1}{2}\partial_\mu\eta\partial^\mu\eta - \frac{3}{2}\sinh^2(\eta/\sqrt{3})(\partial\theta + 2L^{-1}A/\sqrt{3})^2 - V \right], \tag{2.33}$$

with

$$V = -\frac{6}{L^2} \left(1 + \cosh^2(\eta/\sqrt{3}) \right). \tag{2.34}$$

2.3 Non-supersymmetric truncation II: keeping scalars in the 20_A

Next we consider a different truncation from the supersymmetric theory of section 2.1, where we retain $SL(2)$ scalars from the $(\mathbf{2}, 20_A)^+$ rather than the $(\mathbf{1}, 20_S)$ of $SO(2) \times SO(6)$. This truncation was considered in [19, 20]. It can be defined by keeping the invariant sector under a \mathbb{Z}_4 symmetry consisting of (2.6) augmented by the $SO(2) \times SO(6)$ element

$$\left(\begin{array}{cc} & 1 \\ -1 & \end{array} \right) \otimes \left(\begin{array}{ccc} 1 & & \\ -1 & & \\ & 1 & \\ & -1 & \\ & & 1 \\ & & -1 \end{array} \right). \tag{2.35}$$

Surprisingly, this truncation also leads to the scalar manifold (2.12). The origin of the $SO(1,1)^2$ scalars coincides with that of the previous section: these are still given by the scalars $\varphi_{1,2}$ and have masses $m^2 L^2 = -4$. In contrast, in this case the $SL(2)^4$ scalars come from the $(\mathbf{2}, \mathbf{20}_A)^+$ and therefore have masses equal to -3 . Therefore this truncation can be seen as being orthogonal to the one previously considered, at least in the scalar sector. As for the vectors, the above element retains the same three $U(1)$'s as in the previous two truncations.

The bosonic part of the resulting Lagrangian [20] formally looks like the bosonic Lagrangian (2.16) of sector I modulo the following points:

- As pointed out before, the two $SO(1,1)$ dilatons $\varphi_{1,2}$ are common in both truncations, while the $(SL(2)/SO(2))^4$ scalars are completely orthogonal. For that reason we will refer to the charged scalars of sector II as $\tilde{\eta}_a$.
- The charges of the four charged scalars w.r.t. the $U(1)^3$ gauge group in this case are given by

$$q_{1i} = (1, 1, -1), \quad q_{2i} = (1, -1, 1), \quad q_{3i} = (-1, 1, 1), \quad q_{4i} = (1, 1, 1). \quad (2.36)$$

Note that the Hessian of the scalar potential (2.20) in this case indeed gives rise to masses at -3 , corresponding to the upper bound on scalar masses that allow for two different boundary conditions.

Similar to the previous case, we will now consider the simplest Ansätze within this sector which are consistent with configurations of one single charge, two equal charges and three equal charges.

One single charge (II).

- (i) We consider configurations with a single charge, corresponding to $A_i = (0, 0, A)$. In this case it is consistent to set

$$\tilde{\eta}_a = (\eta, 0, 0, \eta)/\sqrt{2}, \quad \tilde{\theta}_a = (-\theta, 0, 0, \theta), \quad \varphi_2 = 0. \quad (2.37)$$

The Lagrangian takes the form

$$\mathcal{L} = \sqrt{-g} \left[R - 3(\partial\varphi_1)^2 - \frac{1}{2}(\partial\eta)^2 - \sinh^2(\eta/\sqrt{2})(\partial\tilde{\theta} + L^{-1}A)^2 - \frac{1}{4}e^{4\varphi_1}F_{\mu\nu}F^{\mu\nu} - V \right], \quad (2.38)$$

where

$$V = \frac{e^{-4\varphi_1}}{L^2} \left(\cosh^2(\eta/\sqrt{2}) - 1 - 4e^{6\varphi_1} - 8e^{3\varphi_1} \cosh(\eta/\sqrt{2}) \right). \quad (2.39)$$

- (ii) We consider the same configuration of gauge vectors, but now we set

$$\tilde{\eta}_a = (\eta, \eta, \eta, \eta)/2, \quad \tilde{\theta}_a = (-\theta, \theta, \theta, \theta), \quad \varphi_2 = 0. \quad (2.40)$$

The Lagrangian takes the form

$$\mathcal{L} = \sqrt{-g} \left[R - 3(\partial\varphi_1)^2 - \frac{1}{2}(\partial\eta)^2 - 2 \sinh^2 \left(\frac{1}{2}\eta \right) (\partial\theta + L^{-1}A)^2 - \frac{1}{4}e^{4\varphi_1} F_{\mu\nu} F^{\mu\nu} - V \right], \quad (2.41)$$

where

$$V = -\frac{2e^{-4\varphi_1}}{L^2} \left((4e^{3\varphi_1} - 1) \cosh^2 \left(\frac{1}{2}\eta \right) + 1 + 2e^{6\varphi_1} \right). \quad (2.42)$$

Two equal charges (II). We now consider configurations with gauge vectors

$$A_i = (A, A, 0)/\sqrt{2}. \quad (2.43)$$

Similar to the previous case, it is consistent to set the scalars to the values (2.37) except for $\tilde{\theta}_a = (\theta, 0, 0, \theta)$. The Lagrangian takes the form

$$\mathcal{L} = \sqrt{-g} \left[R - 3(\partial\varphi_1)^2 - \frac{1}{2}(\partial\eta)^2 - \sinh^2 (\eta/\sqrt{2}) (\partial\theta + \sqrt{2}L^{-1}A)^2 - \frac{1}{4}e^{-2\varphi_1} F_{\mu\nu} F^{\mu\nu} - V \right], \quad (2.44)$$

with V given by (2.39).

Three equal charges (II). Finally we consider the truncation with three equal charges. In this case it is consistent to set

$$\tilde{\eta}_a = (0, 0, 0, \eta), \quad \tilde{\theta}_a = (0, 0, 0, \theta), \quad \varphi_1 = \varphi_2 = 0. \quad (2.45)$$

The remaining scalar η_4 has the highest charge under this diagonal $U(1)$. The resulting Lagrangian is given by

$$\mathcal{L} = \sqrt{-g} \left[R - \frac{1}{2}(\partial\eta)^2 - \frac{1}{2} \sinh^2(\eta) (\partial\theta + L^{-1}\sqrt{3}A)^2 - \frac{1}{4}F_{\mu\nu} F^{\mu\nu} - V \right], \quad (2.46)$$

with

$$V = \frac{3}{L^2} \cosh^2 \left(\frac{1}{2}\eta \right) (\cosh(\eta) - 5). \quad (2.47)$$

We recognize the Lagrangian obtained in [6] by a consistent truncation from IIB theory based on D3-branes at the tip of a Calabi-Yau cone. We can see explicitly how that model (in the case of S^5) also arises from a consistent truncation of $\mathcal{N} = 8$ supergravity. In short, it corresponds to modding out by a \mathbb{Z}_4 symmetry and moreover retaining a diagonal $U(1)$ gauge field and a single complex scalar. The emergence of the model of [6] in the present context is expected. The reason is that the complex scalar field in [6] lives within an $\mathcal{N} = 1$ supergravity model and is the scalar which is charged under the $U(1)$ of R -symmetry; in the present context, the complex scalar field η_4 is the only one that is charged under the $U(1)$ of R -symmetry — the diagonal subgroup of the three $U(1)$'s that we are retaining in the truncation — and at the same time neutral under the rest of this gauge group.

This concludes our discussion of the different truncations of both non-supersymmetric sectors. In table 1 we summarize masses and charges for the two different truncations. We have always normalised the phase θ in such a way the first term in the expansion at small

# Equal charges	Sector I		Sector II	
	qL	m^2L^2	qL	m^2L^2
One	2	-4	1	-3
Two	$\sqrt{2}$	-4	$\sqrt{2}$	-3
Three	$2/\sqrt{3}$	-4	$\sqrt{3}$	-3

Table 1. The charges and masses of the charged scalars of sectors I and II for different configurations of black hole charges.

η has coefficient 1/2, i.e. $\frac{1}{2}\eta^2(\partial\theta)^2 + \dots$. This leads to a universal normalisation for the term $\frac{1}{2}q^2\eta^2A^2$ coming from $(\partial\theta + qA)^2$. Note that the charge of the scalar field η of sector I decreases as one includes more BH charges, while in sector II this is exactly opposite: the charge of $\tilde{\eta}$ increases with the BH charges. The masses are common in both sectors. In the next section we will discuss the interplay between sectors I and II in the context of condensed matter applications.

3 Thermodynamics and condensed matter applications

The truncations obtained in the previous section can be described in terms of the following general Lagrangian

$$\mathcal{L} = \sqrt{-g} \left(R - 3(\partial\varphi)^2 - \frac{1}{4}G(\varphi)F^{\mu\nu}F_{\mu\nu} - \frac{1}{2}\partial_\mu\eta\partial^\mu\eta - \frac{1}{2}J(\eta)A_\mu A^\mu - V(\eta, \varphi) \right), \quad (3.1)$$

where the coupling functions $G(\varphi), J(\eta)$ and the potential $V(\eta, \varphi)$ determine the specific truncation. In the one-charge and two-equal charge cases, the dilaton φ cannot be decoupled and therefore one needs to investigate dilatonic black hole solutions. In the three-equal charge case, the Lagrangian (2.33) and (2.46) are obtained upon setting $\varphi = 0$ in (3.1) and choosing the appropriate $J(\eta)$ and $V(\eta)$. The above Lagrangian provides a useful setup to package the dynamics of the various models. The field η in (3.1) will denote the normalized charged field in either sectors I or II. We will return to the original notation $\tilde{\eta}$ for the sector II charged scalar whenever this distinction is relevant.

3.1 Equations of motion and asymptotic behavior

Given the comprehensive setup (3.1), our ansatz for the metric and the gauge field is

$$ds^2 = -g(r)e^{-\chi(r)}dt^2 + \frac{dr^2}{g(r)} + r^2(dx^2 + dy^2 + dz^2), \quad A = \Phi(r)dt, \quad (3.2)$$

with the scalars fields being functions only of the radial coordinate. The equations of motion derived from the Lagrangian (3.1) are then given by

$$\frac{3}{2}\chi' + 3r\varphi'^2 + \frac{r}{2}\eta'^2 + \frac{r}{2g^2}e^\chi J(\eta)\Phi^2 = 0, \quad (3.3)$$

$$3\left(\frac{g'}{gr} + \frac{2}{r^2}\right) + \frac{1}{2}\left(\eta'^2 + J(\eta)e^\chi\frac{\Phi^2}{g^2}\right) + 3\varphi'^2 + \frac{e^\chi}{2g}G(\varphi)\Phi'^2 + \frac{V(\eta, \varphi)}{g} = 0, \quad (3.4)$$

$$\Phi'' + \Phi' \left(\frac{3}{r} + \frac{\chi'}{2} + \frac{\partial_\varphi G}{G} \varphi' \right) - \frac{J(\eta)}{gG(\varphi)} \Phi = 0, \quad (3.5)$$

$$\varphi'' + \varphi' \left(\frac{3}{r} - \frac{\chi'}{2} + \frac{g'}{g} \right) + \frac{e^\chi}{12g} \partial_\varphi G \Phi'^2 - \frac{\partial_\varphi V(\eta, \varphi)}{6g} = 0, \quad (3.6)$$

$$\eta'' + \eta' \left(\frac{3}{r} - \frac{\chi'}{2} + \frac{g'}{g} \right) + \frac{e^\chi}{2g^2} \partial_\eta J \Phi^2 - \frac{\partial_\eta V(\eta, \varphi)}{g} = 0. \quad (3.7)$$

We are interested in black hole configurations that have a regular event horizon and appropriate boundary conditions at infinity. It is easy to check that all truncations given above possess an AdS vacuum for which $\eta = 0$, $\varphi = 0$, $\Phi = 0$. Therefore, in order to apply the AdS/CFT dictionary we will look for asymptotically AdS black holes and require that $g \approx r^2/L^2$ as r goes to infinity. Linearizing the equations around this AdS vacuum we find that a generic solution has the following asymptotic behavior

$$\varphi = \frac{O_\varphi}{r^2} + \frac{C_\varphi}{r^2} \log r + \dots \quad (3.8)$$

$$\Phi = \mu - \frac{\rho}{r^2} + \dots \quad (3.9)$$

$$\chi = \chi_\infty + \dots \quad (3.10)$$

$$e^{-\chi} g(r) = e^{-\chi_\infty} \left(\frac{r^2}{L^2} - \frac{\epsilon}{r^2} + \dots \right), \quad (3.11)$$

where the dots stand for terms of higher order in the expansion in powers of $1/r$. The asymptotic of the dilaton is universal in all models; indeed in all cases φ is a scalar field in AdS with $m^2 L^2 = -4$. On the other hand, the charged scalars have different masses in sector I and in sector II. In sector I these charged scalars have $m^2 L^2 = -4$ and they have the same asymptotic behavior as the dilaton,

$$\eta = \frac{O_\eta}{r^2} + \frac{C_\eta}{r^2} \log r + \dots \quad (3.12)$$

while in the sector II these correspond to scalars with $m^2 L^2 = -3$, therefore

$$\tilde{\eta} = \frac{C_\eta}{r} + \frac{O_{\tilde{\eta}}}{r^3} + \dots \quad (3.13)$$

The location of the horizon $r = r_h$ is defined by the simple zero of g lying at the largest r . The Hawking temperature of the black hole can then be calculated as usual by the formula,

$$T_{\text{Hawk}} = \frac{1}{4\pi} g'(r) e^{-\chi(r)/2} \Big|_{r=r_h}. \quad (3.14)$$

The value of $g'(r_h)$ is determined from the first order equation (3.4), in particular by the combination

$$g'(r_h) = -r_h \left[V(\eta_h, \varphi_h) + \frac{1}{2} e^{\chi_h} G(\varphi_h) E_h^2 \right], \quad (3.15)$$

where we have defined the parameters

$$\eta(r_h) = \eta_h, \quad \varphi(r_h) = \varphi_h, \quad \Phi'(r_h) = E_h, \quad \text{and} \quad \chi(r_h) = \chi_h. \quad (3.16)$$

Finding an analytic solution of the complete system (3.3)–(3.7) is, in general, a very difficult task. The particular case $\eta \equiv 0$ is much simpler because the Maxwell equations can then be integrated and the solution can be substituted into the other equations. In particular, for the non-supersymmetric truncations that we are considering, analytic solutions do exist and are described by the “STU” black hole [24]. Explicit formulas are given in the next section. At any rate, the complete system (3.3)–(3.7), including the back-reaction of the geometry, can be generically solved using numerical methods, including the case $\eta \equiv 0$.

The strategy is as follows. We specify the initial data at the horizon in such a way that the Cauchy problem (3.3)–(3.7) is well posed and then numerically integrate up to infinity. A priori, we have nine parameters to deal with: the location of the horizon r_h and eight initial conditions for the equations of motions. The requirement $g(r_h) = 0$ constrains some of them. In particular, the consistency of equations (3.6) and (3.7) fixes the values of $\eta'(r_h)$ and $\varphi'(r_h)$ to satisfy the relations

$$\begin{aligned} g'(r_h)\eta'(r_h) &= \partial_\eta V(\eta_h, \varphi_h), \\ g'(r_h)\varphi'(r_h) &= \frac{1}{6}\partial_\varphi V(\eta_h, \varphi_h) - \frac{1}{12}e^{\chi_h}\partial_\varphi G(\varphi_h) E_h^2, \end{aligned} \tag{3.17}$$

whereas the condition $\Phi(r_h) = 0$ is needed to ensure that the gauge field is well defined at the horizon. Therefore, out of the nine parameters we started with, only the ones in the set $\{r_h, \eta_h, \varphi_h, E_h, \chi_h\}$ are independent. This set can be further reduced. In order to do so, we note the following two scaling symmetries,

$$r \rightarrow ar, \quad (t, \vec{x}) \rightarrow a^{-1}(t, \vec{x}), \quad g \rightarrow a^2g, \quad \Phi \rightarrow a\Phi, \tag{3.18}$$

and

$$e^\chi \rightarrow a^2 e^\chi, \quad t \rightarrow at, \quad \Phi \rightarrow a^{-1}\Phi, \tag{3.19}$$

which leave invariant the metric, the gauge field and the equations of motion. These two symmetries can be used to set $r_h = 1$ and $\chi_\infty = 0$. Thus, each black hole solution obtained by integrating the equations of motion from the horizon is characterized in terms of the three horizon parameters η_h, φ_h, E_h .

Finally, the relevant configurations will be the ones with the asymptotic constraints $C_\varphi = 0$ and $C_\eta = 0$ (or $C_{\tilde{\eta}} = 0$). These two additional conditions leave a one-parameter family of hairy black hole solutions, where the parameter characterizing the solution can be taken to be the temperature. These configurations represent the physical systems that we are going to describe.

It should be noted that in all truncations it is consistent to look for solutions with vanishing η (which, in particular, implies $C_\eta = 0$ or $C_{\tilde{\eta}} = 0$). Therefore there will be two kinds of interesting configurations: hairy black holes, with a non-trivial η turned on, and “bald” black holes with $\eta \equiv 0$.

The field theory interpretation of these two types of solutions follows from the standard AdS/CFT dictionary, that we now briefly review. This asserts that each field in the gravitational action is dual to a certain operator in the conformal field theory. In five dimensions the correspondence between a scalar field in the bulk with some operator in

the CFT of conformal dimension Δ is given by the relation $m^2 L^2 = \Delta(\Delta - 4)$. Consider, for concreteness, the scalar $\tilde{\eta}$ of sector II. Its mass is given by $m^2 L^2 = -3$, therefore there are two solutions: $\Delta_- = 1$ and $\Delta_+ = 3$. This fact shows up in the powers of r arising from the behavior of the field at infinity,

$$\tilde{\eta} = \frac{C_\eta}{r} + \frac{O_{\tilde{\eta}}}{r^3} + \dots$$

Among the two terms, only the mode associated with $\Delta_+ = 3$ is normalizable; thus, one interprets $O_{\tilde{\eta}}$ as the condensate associated with an operator of dimension 3 in the presence of the external source C_η . Demanding that the solution has $C_{\tilde{\eta}} = 0$ implies that the field theory is not sourced by this operator; a non-zero value of $O_{\tilde{\eta}}$ then implies a vacuum expectation value for the CFT operator of dimension 3. In other words the AdS/CFT correspondence implies that a hairy black hole with appropriate boundary conditions is dual to a condensed phase of the system, while a “bald” black hole represents an uncondensed phase. Moreover, because $\tilde{\eta}$ is charged under a U(1) symmetry, its dual operator will carry the same charge under this symmetry and a vacuum expectation value of this operator then implies U(1) spontaneous symmetry breaking.

For the charged scalar η of sector I there is a unique solution to the mass/dimension relation. This scalar is dual to a charged operator of dimension $\Delta = 2$. The dilaton is dual to an operator of the same dimension, but neutral under the U(1) symmetry. In these cases, when $m^2 = m_{BF}^2$, a logarithmic branch appears in the asymptotic. Such a branch will necessarily introduce an instability unless it is treated as a source [25]. This means that one must set $C_\eta = 0$ and the interpretation then follows as for the $\tilde{\eta}$ field.

3.2 The uncondensed phase

In the maximal SO(6) gauged supergravity, black hole solutions carrying arbitrary charges with respect to three different U(1) symmetries are described by the “STU” black hole. The present truncations contain these black holes since they maintain the U(1) \times U(1) \times U(1) gauge symmetry and thus the three relevant vector fields. Indeed, keeping only the real scalar fields in sector I and sector II, we recover the Lagrangian of [24],

$$\mathcal{L} = \sqrt{-g} \left(R - \sum_{i=1}^3 X_i^{-2} \left[\frac{1}{2} (\partial X_i)^2 + \frac{1}{4} F^{i,\mu\nu} F_{\mu\nu}^i \right] \right), \quad X_1 X_2 X_3 = 1, \quad (3.20)$$

for which the following analytic solution was found:

$$\begin{aligned} ds^2 &= -f H^{-2/3} dt^2 + H^{1/3} f^{-1} dr^2 + H^{1/3} \frac{r^2}{L^2} d\vec{x}^2, \\ A_i &= \left(\frac{Q_i \sqrt{m}}{r_h^2 + Q_i^2} - \frac{Q_i \sqrt{m}}{r^2 + Q_i^2} \right) dt, \quad X_i = H_i^{-1} H^{1/3}, \\ f &= \frac{r^2}{L^2} H - \frac{m}{r^2}, \quad H = H_1 H_2 H_3, \quad H_i = 1 + \frac{Q_i^2}{r^2}. \end{aligned} \quad (3.21)$$

The charged scalar fields are vanishing, therefore these solutions describe the uncondensed phase for our models. The solution (3.21) can be written in the form (3.2) by a coordinate

transformation on the metric: $r^2 H^{1/3} \leftrightarrow r^2$. It is easy to check that the solution (3.21) satisfies the equations of motion (3.3)–(3.7). For the case of three equal charges, the coordinate transformation simply becomes $r^2 + Q^2 \leftrightarrow r^2$ and one recovers the standard form of the AdS Reissner-Nördstrom black hole.

It is useful to review the relevant quantities for the field theory thermodynamics in the uncondensed phase (see [10, 26] for recent discussions). The energy, entropy and charge densities are

$$\hat{\epsilon} = \frac{3m}{8\pi^2 L^6}, \quad \hat{s} = \frac{r_h \sqrt{m}}{2\pi L^5}, \quad \hat{\rho}_i = \frac{Q_i \sqrt{m}}{4\pi^2 L^5}, \quad (3.22)$$

where the quantities m and Q , in general, can be read from the asymptotics of the solution. The position of the horizon is found by solving the equation $f = 0$. In general there will be three roots. We use the notation r_h^2, r_1^2 and $-r_0^2$ for these roots indicating with r_h^2 the greatest (real) one. The event horizon is then located at r_h and the temperature is given by

$$T_H = \frac{(r_h^2 + r_0^2)(r_h^2 - r_1^2)}{2\pi L^2 \sqrt{(r_h^2 + Q_1^2)(r_h^2 + Q_2^2)(r_h^2 + Q_3^2)}}. \quad (3.23)$$

While the STU black hole represents an exact solution for arbitrary choice of the (Q_1, Q_2, Q_3) , we will be interested in three particular cases: one charge, two-equal charges and three-equal charges. In what follows we show how the uncondensed black holes solutions of sector I and sector II are obtained from the STU black hole solution in each case. Finally we observe that having the analytic solution for the uncondensed phases allows us to check the thermodynamical quantities provided by the numerics.

Black hole with one single charge. The first case of interest is when $Q_i = (0, 0, Q)$. The solution (3.21) is characterized by the two functions

$$H = \left(1 + \frac{Q^2}{r^2}\right) \quad \text{and} \quad A_3 = \left(\frac{Q_i \sqrt{m}}{r_h^2 + Q_i^2} - \frac{Q_i \sqrt{m}}{r^2 + Q_i^2}\right) dt. \quad (3.24)$$

This black hole represents the uncondensed solution for the Lagrangians (2.23) and (2.38), in which we set $A_3 = A$. The charge density associated to the U(1) gauge group is $\hat{\rho} = Q\sqrt{m}/(4\pi^2 L^5)$.

Black hole with two equal charges. In this case two gauge fields are identified and we set $Q_i = (Q, Q, 0)$ in (3.21). The black hole is then specified by the functions

$$H = \left(1 + \frac{Q^2}{r^2}\right)^2 \quad \text{and} \quad A_1 = A_2. \quad (3.25)$$

This is a solution to the equations of motion of the Lagrangians (2.28) and (2.44) with the identifications

$$A_1 = A_2 = A/\sqrt{2}, \quad X_1 = X_2 = e^{\varphi_1}. \quad (3.26)$$

The charge density associated to the diagonal U(1) is therefore $\hat{\rho} = Q\sqrt{2m}/(4\pi^2 L^5)$.

Black hole with three equal charges. This is the case when $Q_i = (Q, Q, Q)$ and the three gauge fields are identified. Now we have

$$H = \left(1 + \frac{Q^2}{r^2}\right)^3 \quad \text{and} \quad A_1 = A_2 = A_3. \quad (3.27)$$

This represents a solution for the uncondensed phase for the models (2.33) and (2.46) with the identifications

$$A_1 = A_2 = A_3 = A/\sqrt{3}, \quad X_1 = X_2 = X_3 = 1. \quad (3.28)$$

The charge density then is given by $\hat{\rho} = Q\sqrt{3m}/(4\pi^2 L^5)$.

Let us now discuss some basic aspects of the thermodynamics at fixed $\hat{\rho}$. The Hawking temperature is given by the formula

$$T = \frac{1}{4\pi} H(r_h)^{-1/2} f'(r_h). \quad (3.29)$$

The horizon equation, $f = 0$, gives

$$m = \frac{r_h^4}{L^2} H(r_h) \quad \text{where} \quad H = \left(1 + \frac{Q^2}{r_h^2}\right)^\alpha \quad \alpha = 1, 2, 3, \quad (3.30)$$

with α counting the number of equal charges. It is convenient way to rescale $r \rightarrow r_h \tilde{r}$, so the horizon is at $\tilde{r} = 1$. We also introduce new parameters \tilde{m} and \tilde{Q} by the rescaling

$$m = r_h^4 \tilde{m}, \quad Q = r_h \tilde{Q}. \quad (3.31)$$

Then the relation (3.30) simplifies, reducing to the formula $\tilde{m} = \tilde{H}/L^2 = (1 + \tilde{Q}^2)/L^2$. Using these relations, the temperature (3.29) becomes

$$T = \frac{r_h}{2\pi L^2} (1 + \tilde{Q}^2)^{1-\alpha/2} (2 + (2 - \alpha)\tilde{Q}^2). \quad (3.32)$$

We are interested in the thermodynamics in the fixed $\hat{\rho}$ thermal ensemble. In terms of the rescaled variables, the charge density $\hat{\rho}$ takes the form

$$\hat{\rho} = \frac{\tilde{Q}\sqrt{\alpha\tilde{m}}}{4\pi^2 L^5} r_h^3 = \sqrt{\alpha} \frac{\tilde{Q}(1 + \tilde{Q}^2)^{\alpha/2}}{4\pi^2 L^6} r_h^3 \quad (3.33)$$

Thus, solving for r_h , the temperature becomes

$$T = \frac{\tilde{Q}^{-1/3}}{(2\pi\sqrt{\alpha})^{1/3}} (2 + (2 - \alpha)\tilde{Q}^2) (1 + \tilde{Q}^2)^{\frac{\alpha}{3}-1} \hat{\rho}^{1/3}. \quad (3.34)$$

Since $\hat{\rho}$ is fixed, the temperature T is only a function of the auxiliary variable \tilde{Q} . From this expression one can get a qualitative understanding of the properties enjoyed by the different ensembles. In the one charge case $\alpha = 1$ and T is a strictly positive function that goes to infinity as \tilde{Q} approaches the two limits, $\tilde{Q} \ll 1$ and $\tilde{Q} \gg 1$, therefore the uncondensed phase has a minimum temperature whose significance was extensively discussed in [10]. We shall see that the system can be driven to lower temperatures by condensation; in

other words, there are black hole solutions with hair that reach lower temperatures. In the two-equal charge ensemble one has $\alpha = 2$ and T varies between 0 and infinity. Some aspects of the thermodynamics of these solutions are discussed in [26]. For the purpose of this work, we point out that this black hole reproduces some features of a Fermi-liquid phase and therefore the search for the condensed phase is well motivated. Some of these features are, in particular, the linear specific heat and zero entropy as $T \rightarrow 0$. In the AdS Reissner-Nördstrom black hole, $\alpha = 3$ and \tilde{Q} is restricted to be in the interval $\tilde{Q} \in [0, \sqrt{2}]$. The temperature vanishes as $\tilde{Q} \rightarrow \sqrt{2}$, but the horizon has a finite size resulting in a large value for the entropy at zero temperature.

3.3 The condensed phase

In what follows we will look for a condensed phase in the different sectors. These truncated models could be regarded as holographic setups for AdS/CMT applications on their own right. Nevertheless, the most interesting aspect of top-down constructions, like the present one, is represented by the explicit knowledge of the dual field theory.⁶ We hope that this explicit connection will open the way to the study of novel features of the thermodynamics of large N SYM in the strong coupling regime. On the field theory side, some aspects of its phase diagram have been discussed on S^3 (see e.g. [28, 29]) but so far there have not been many discussions on the spontaneous symmetry breaking of U(1) symmetries.

In principle, in order to understand the thermodynamics of the system at fixed charge densities, one should search for the dominant thermodynamic configuration not only in a given sector but in the full ten-dimensional theory. This would of course be a complicated task already in the context of $D = 5$ maximal SO(6) supergravity, because one should look for the hairy black hole configuration with least free energy among configurations where any of the 42 scalars can be turned on (or, more generally, even one-form or two-form hair as in [10]). Despite this huge number of possibilities, an important glimmer of information is provided by the observation that not all sectors are relevant in order to identify the highest critical temperature at which the first phase transition can occur. Indeed, one can show (see e.g. figure 2 in [6]) that the critical temperature increases with the charge and decreases with the mass of a given mode. This reflects the fact that, for given charge, the dual operators with the minimal dimension are those which should condense first; whereas, for given conformal dimension, operators with highest R-charge should condense first. The competition between charge and dimension will appear in several examples below. Furthermore, in section 4 we will also discuss the possibility of condensation arising from some of the 42 scalars not included in our truncated Lagrangians.

The different sectors studied in this paper also include modes⁷ with lowest possible mass (saturating the BF bound) which are therefore dual to (protected) operators of lowest dimensions. Nonetheless, we shall see that their dual operators do not always condense before other operators of the theory: in some cases there are modes which, while having

⁶Another approach where the dual field theory is known is based on using D-brane probes in string-theory black brane backgrounds (see e.g. [27] and references therein).

⁷Some special features in the conductivity arise for scalars saturating the BF bound, see [25].

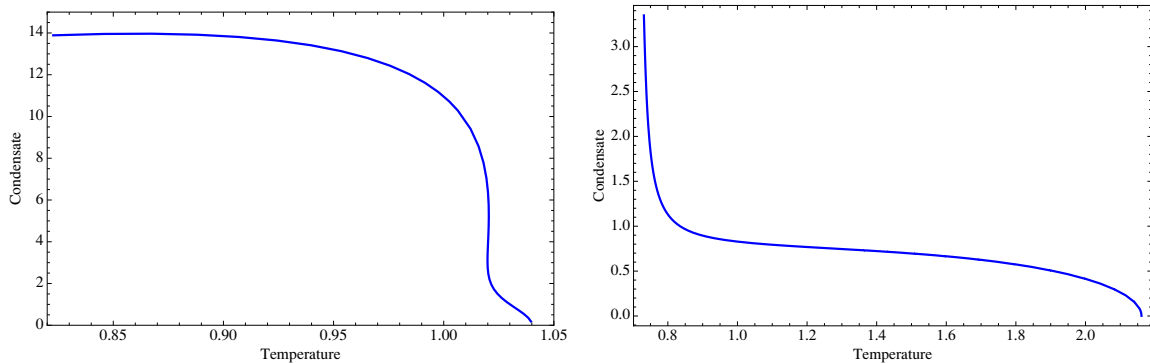


Figure 2. (a) Condensate as a function of the temperature for the one charge model in sector I ($q = 2$). (b) Similar plot, for a scalar with the same Lagrangian but charge $q = 16$ instead of $q = 2$.

higher mass, they have a sufficiently larger charge to trigger condensation at a higher critical temperature.

In what follows we shall numerically exhibit the competition between sectors I and II. All the plots shown in this section are obtained in the canonical ensemble at fixed charge density $\hat{\rho} = 1$. For the uncondensed black holes solutions we refer to the formulas already given case by case in section 3.2. We recall that $\hat{\rho} = \rho/(4\pi^2 L^5)$, where ρ is the asymptotic value that we read from the expansion (3.9) as usual. Similarly for the energy \hat{e} . In the gauge (3.2) the entropy and then the free energy density are given by

$$\hat{s} = \frac{r_h^3}{2\pi L^6}, \quad \hat{f} = \hat{e} - T\hat{s}. \quad (3.35)$$

Condensation in single-charge thermal ensemble.

- *Condensation in Sector I*

Consider the Lagrangian (2.23), (2.24). It describes a complex scalar of charge $qL = 2$ and $m^2 L^2 = -4$ and a real scalar φ of mass $m^2 L^2 = -4$. We have numerically computed the critical curve representing the order parameter as a function of the temperature, including backreaction (see figure 2a). We find that there is a critical temperature $T_c \cong 1.04$ below which a hairy black hole appears. This curve presents some unusual features.

At some lower temperature $T_1 \cong 1.02$ the second derivative of the condensate with respect to the temperature changes sign. This could be an indication of a new phase transition as it implies that the fourth order term of the free energy in terms of the condensate must have a strong temperature dependence near T_1 . In Landau-Ginzburg theory, it is normally assumed that the fourth order coefficient is not strongly temperature dependent around the critical point T_c , but it is in principle possible that there is some new temperature scale below T_c where this coefficient also starts to change. In particular, if it goes to zero or becomes negative then this term no longer stabilizes the condensate, and higher order coefficients (if present) become important.

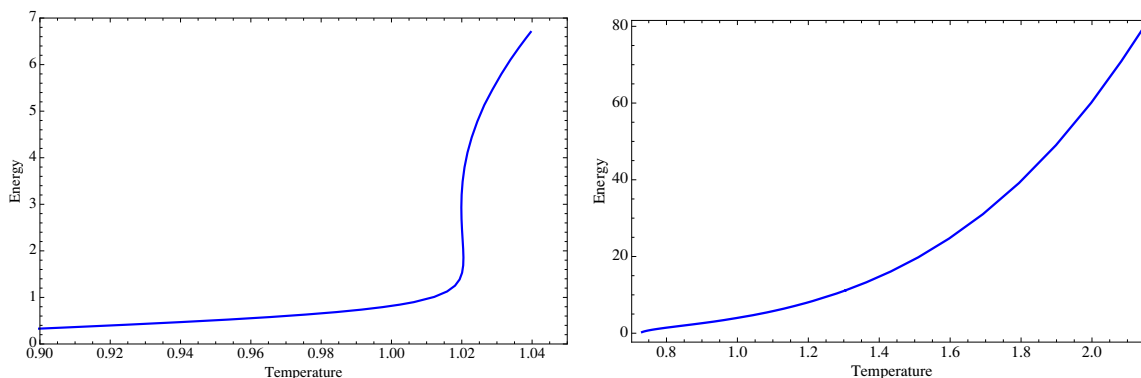


Figure 3. (a) Energy as a function of the temperature for the one-charge model in sector I. (b) Similar plot, for a scalar with the same Lagrangian but charge $q = 16$ instead of $q = 2$.

Because of this, the behavior of the condensate can change and indeed can blow up at this new temperature scale, as indeed occurs in the present model.

This behavior can be seen with more clarity by considering a scalar field with the same Lagrangian but larger charge q . In figure 2b we have plotted the condensate as a function of the temperature for $q = 16$. One can see that, near T_c , the behavior is very much like in second order phase transitions of ordinary mean field theories. Then, at some lower temperature $T_1^{(q=16)} \cong 0.75$ the second derivative of the order parameter changes sign and the order parameter then increases very rapidly. For $q \rightarrow \infty$, one reaches the probe limit where the dilaton and metric are decoupled, and one can study the system in terms of a charged scalar in the AdS Schwarzschild black hole background. In this limit, $T_1^{(q)} \rightarrow 0$, i.e. there is no change of sign in the second derivative of the condensate curve, which has the usual mean field shape all the way down to $T = 0$.

Returning to the $q = 2$ case, on the gravity side, as the temperature is slightly lowered below T_1 , the black hole gets drastically reduced to a very small size, i.e. r_h becomes very small, and the internal energy $\hat{\epsilon}$ of the system goes down abruptly to almost zero value. This is shown in figure 3a. As a result, the specific heat has a very large peak near T_1 . In the $q = 16$ case, the internal energy also goes to almost zero value, but at a slower rate (see figure 3b). One common feature which is present in both cases is that the condensate is catapulted to large values at temperatures just below the inflexion point in the critical curves of figure 2a,b. Another common feature is the emergence of a hairy black hole with very small entropy and energy at finite T . Because these quantities are obtained from the asymptotic behavior of the solution, when they get small, it becomes difficult to determine them with sufficient accuracy. The numerical results seem to indicate that energy and entropy remain small and smoothly decrease as the temperature is lowered from T_1 to small values.

We have computed the free energy of both the $q = 2$ and the $q = 16$ hairy black holes and found that these are indeed lower than the free energy of the dilatonic bald black hole describing the uncondensed phase. Therefore it dominates the thermodynamics

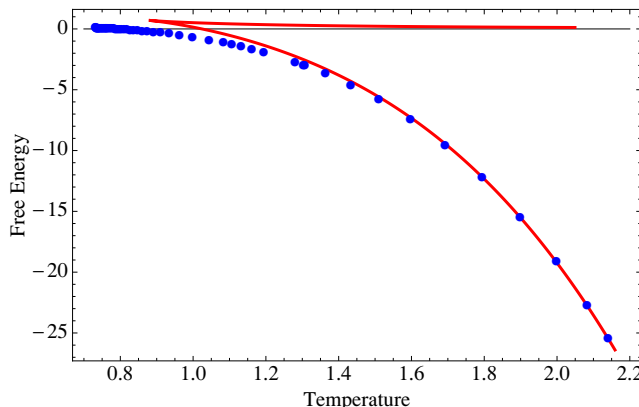


Figure 4. Free energy for the system described by the Lagrangian (2.23), (2.24) with $q = 16$ (one charge ensemble, sector I). The solid lines represent the two branches of the bald dilatonic black hole solution. The dotted line represents the free energy for the hairy black hole solution.

at $T < T_c$. Figure 4 shows the free energy for the $q = 16$ model, which is easier to visualize. One can see that there is an accumulation of points approaching zero free energy, at temperatures just below $T_1^{(q=16)} \cong 0.75$. The deep physical reason for this is mysterious to us, but it is a consequence of the fact that the internal energy $\hat{\epsilon}$ and entropy \hat{s} get very small near that point.

It is worth noting that the hairy black hole solutions reach temperatures which are lower than the minimum temperature of the dilatonic black hole describing the uncondensed phase. For the case $q = 2$, the numerical results show that the hairy black hole solution exist up to a new minimum temperature around 0.22, where it joins an unstable branch coming from higher temperatures. Nevertheless, already at temperatures lower than $T_1 \cong 1.02$ it is far from clear that the solution is describing reliable physics. It could also be that there is another condensate at the temperature T_1 , i.e. there is some other state with lower free energy that arises exactly at this new temperature.

- *Condensation in Sector II*

Model (i): the relevant Lagrangian is (2.38), (2.39). It contains a complex scalar field of charge $qL = 1$ and $m^2 L^2 = -3$, plus a real scalar φ . Note that it has a lower charge and larger mass than the previous sector I scalar. Thus one can expect that, if it condenses, it will be at a lower critical temperature.

By solving the equations numerically we find that for that particular value of the charge there is no condensation, i.e. there is no value for the temperature at which solutions exist with the required boundary conditions. The origin of this lack of condensation is not only the fact that the scalar has lower charge and larger mass, but it can also be traced back to the coupling between η and the dilaton φ through the potential $V(\eta, \varphi)$. In the vicinity of the critical temperature, η is small everywhere and one can study the emergence of the hairy black hole by the linearized equation

for η in the bald (dilaton) black hole background. It has the schematic form:

$$\eta'' + F(r)\eta' + q^2\eta\Phi_{\text{uncond.}}^2 - m^2\eta\tilde{V}(\varphi_{\text{uncond.}}) = 0. \quad (3.36)$$

At infinity the dilaton goes to zero and we have fixed the normalization of \tilde{V} to be $\tilde{V}(0) = 1$. In non-dilatonic models, the critical temperature depends only on qL and m^2L^2 . However, in dilatonic models we find that the critical temperature depends also on the specific form of $\tilde{V}(\varphi_{\text{uncond.}})$. In particular, one can find examples of $\tilde{V}(\varphi_{\text{uncond.}})$ which give rise to condensation for a scalar η with the same m^2 and q as in this model.

Model (ii): now the relevant Lagrangian is (2.41), (2.42). As in the previous case, we have a complex scalar field of charge $qL = 1$ and $m^2L^2 = -3$, plus a real scalar φ . We again find that there is no condensation in this system for similar reasons as in the model (i).

Condensation in two equal-charge thermal ensemble.

- *Condensation in Sector I*

The relevant Lagrangian is given in (2.28), (2.29). It contains a complex scalar of charge $qL = \sqrt{2}$ and $m^2L^2 = -4$ and a real scalar φ . Here we find that there exists a family of hairy black holes parametrized by the temperature for any⁸ $T > T_c$. On the face of it, this appears to be surprising, since one expects that the condensed phase appears at low temperatures, not at high temperatures. But for condensation to actually take place, it is necessary that the free energy of this hairy black hole configuration be less than the free energy of the two-equal charge STU black hole describing the uncondensed phase. We find that the free energy of such hairy black holes is at all temperatures $T > T_c$ greater than the free energy of this STU black hole. Therefore these hairy black holes represent unstable branches that do not contribute to the thermodynamics. We will refer to this phenomenon of a thermodynamically subdominant condensate at $T > T_c$ as *retrograde condensation*.⁹

The results are shown in figure 5a,b. One may also wonder if the critical curve could turn back at even higher temperatures not displayed in figure 5a. We do not expect that this will happen, because figure 5a includes temperatures that are high enough to be above any dimensionful scale of the problem and the curve seems to have already reached a well defined asymptotic behavior.

- *Condensation in Sector II*

We now consider the Lagrangian (2.44), with the potential given in (2.39). The complex scalar now has $qL = \sqrt{2}$ and $m^2L^2 = -3$. Note that it has lower charge and higher mass than the sector I scalar. We find that there is no condensation in this model.

⁸A similar phenomenon was found in four-dimensional models for certain black holes with AdS₄ asymptotic in [30, 31].

⁹The term “retrograde condensation” was first used by Kuenen in 1892 [32] to describe the behavior of a binary mixture during isothermal compression above the critical temperature of the mixture (a discussion can be found in [33]). Such a system also displays the phenomenon of a subdominant condensate in some temperature range.

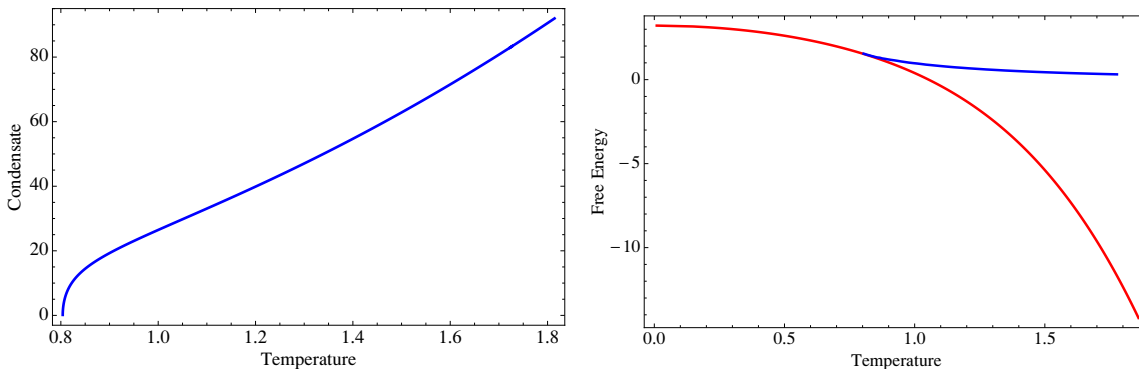


Figure 5. (a) Condensate as a function of the temperature in the two-equal charge thermal ensemble, sector I model. (b) Free energy for the same model. The upper (blue) branch is the free energy of the hairy black hole solution, while the lower red branch is the one for the dilatonic STU black hole with $Q_1 = Q_2, Q_3 = 0$.

Condensation in three equal-charge thermal ensemble. In this case the uncondensed phase is described by the Reissner-Nordström black hole. The real scalars φ_1, φ_2 are set to zero, which greatly simplifies the analysis. The Lagrangians (2.33), (2.46) are examples of the general models introduced in [34, 35],

$$\mathcal{L} = \sqrt{-g} \left[R - \frac{1}{4}G(\eta) F^{\mu\nu} F_{\mu\nu} - \frac{1}{2}\partial_\mu\eta\partial^\mu\eta - \frac{1}{2}J(\eta)A_\mu A^\mu + \frac{12}{L^2}U(\eta) \right], \quad (3.37)$$

with the identification

$$G(\eta) = 1, \quad U(\eta) = \frac{1}{2} \left(1 + \cosh^2(\eta/\sqrt{3}) \right), \quad J(\eta) = \frac{4}{L^2} \sinh^2(\eta/\sqrt{3}), \quad (3.38)$$

for sector I, and

$$G(\eta) = 1, \quad U(\eta) = \frac{1}{4} \cosh^2\left(\frac{1}{2}\eta\right) (5 - \cosh(\eta)), \quad J(\eta) = \frac{3}{L^2} \sinh^2(\eta), \quad (3.39)$$

for sector II.

Some relevant features of the condensation can be learned from the properties of the functions $U(\eta)$ and $J(\eta)$. We borrow part of the discussion in [36, 37], originally given for 3+1 dimensional holographic models in the no-backreaction approximation. Near the critical temperature, η is small and one has the expansions

$$J = q^2\eta^2 + j_0\eta^4 + \dots, \quad (3.40)$$

$$V = -\frac{12}{L^2}U = -\frac{12}{L^2} \left(1 - \frac{m^2L^2}{24}\eta^2 - v_0\eta^4 + \dots \right). \quad (3.41)$$

The quartic terms in J and V play an important role. This can be understood by using the following formula, deduced in [36] (see eq. (3.14)), giving the temperature dependence of the order parameter in the vicinity of a second-order transition:

$$1 - \frac{T}{T_c^{(N)}(m^2, q^2)} = \langle O_1 \rangle^2 (A_N(m^2, q^2) + v_0 C_N + j_0 L^2 D_N) + \dots \quad (3.42)$$

Here dots represent terms with higher powers of the order parameter and $A_N, C_N, D_N > 0$ are numerical coefficients. This formula was derived using a series expansion near the horizon, and here N represents the truncation order. When $v_0 = j_0 = 0$, there are no quartic corrections in U and J functions and the formula reproduces the mean field properties of the condensate as a function of the temperature of the HHH model¹⁰ [2]. The curve will be qualitatively the same near T_c as long as the coefficient of $\langle O_1 \rangle^2$ is positive. In particular, this is the case if $v_0, j_0 > 0$. A drastic change occurs when $v_0 C_N + j_0 L^2 D_N < -A_N$. As we shall see, this condition can be met in models where v_0 is sufficiently negative. Then the condensed phase appears at $T > T_c$, instead of $T < T_c$, at least near T_c . What happens next depends on higher order terms in V and J . In particular, in some cases, the critical curve comes back to the lower temperature region $T < T_c$ (examples of this behavior, representing first order phase transitions, are in [34, 35]). This depends on higher order terms in the potential so it requires solving the full system of equations including backreaction, as we do in this paper. In other examples — like in the two-equal charge ensemble discussed above or in the three-equal charge, sector I, discussed below — the hairy black hole branch extends all the way to the region $T > T_c$, and thus displays the phenomenon of retrograde condensation.

- *Condensation in Sector I*

The relevant Lagrangian is given in (2.33).¹¹ It should be noted that, despite the fact that the potential ($V = -12U/L^2$ with U given in (3.38)) is unbounded, the theory around the trivial stationary point is stable due to the fact that in the original $\mathcal{N} = 8$ supergravity theory this point is supersymmetric (which in particular ensures that all scalar fluctuations have masses at or above the BF bound).

We have numerically integrated the equations and found a family of black hole solutions with charged scalar hair with the correct asymptotic. Figure 6a shows $\langle O_1 \rangle$ vs. T . Like in the two equal-charge case for sector I, we again find retrograde condensation: the hairy black hole solution exists for $T > T_c$, instead of $T < T_c$ (with $T_c \cong 0.56$). This is consistent with the fact that the quartic coefficient v_0 in the potential is negative.

This hairy black hole solution again has free energy which is larger than the free energy of the Reissner-Nordström black hole. This is shown in figure 7. Therefore this condensed phase is unstable and not physically relevant. As in the two-charge case, one can also see that the critical curve already reaches some well-defined asymptotic behavior so it is not expected to turn around.

In the context of Landau-Ginzburg theory, as discussed above, this departure from mean field theory is a sign of an unusual temperature dependence of the coefficients

¹⁰Note that the simplest phenomenological model [2] with $V = m^2 \eta^2 / 2$ with $m^2 < 0$ has an unbounded potential. Nevertheless, due to the AdS boundary conditions, the model exhibits a second-order phase transition with standard (mean field type) critical curve and in particular does not present any runaway behavior.

¹¹A study of hairy black hole solutions which are asymptotic to *global* AdS₅ in this sector was carried out in [38].

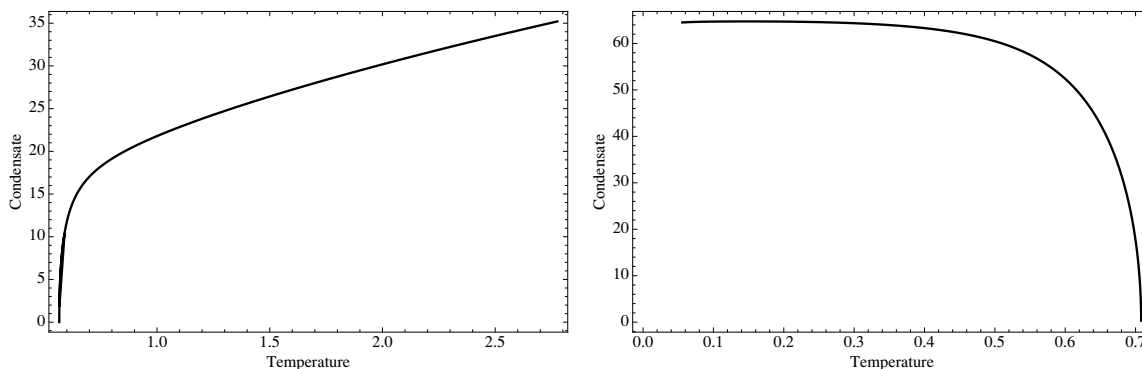


Figure 6. (a) Condensate as a function of the temperature in the three-equal charge thermal ensemble, sector I model. (b) Condensate as a function of the temperature in sector II (representing the same model of [6], now considered in the thermal ensemble at fixed charge density rather than at fixed chemical potential).

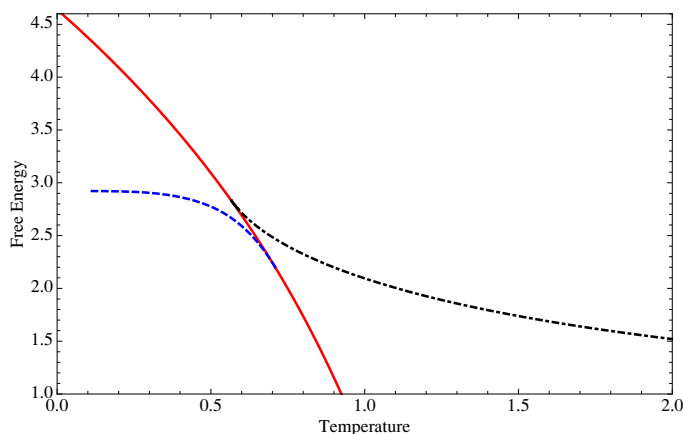


Figure 7. Free energies for the different black hole solutions. The dot-dashed black curve (lying on the $T > T_c$ side, $T_c \cong 0.56$) describes the free energy of the hairy black hole of sector I. The red solid curve (i.e. the one shown in the interval $0 < T < 0.93$, but extending to all T) represents the free energy of the Reissner-Nordström black hole with $Q_1 = Q_2 = Q_3$. Finally, the dashed blue curve lying at $T < 0.70$ describes the free energy of the sector II black hole.

of higher order terms. For example, this behavior might arise if for $T > T_c$ the Landau-Ginzburg potential has a relative maximum at some positive value, which joins the absolute minimum at zero for $T = T_c$. This can be described by a potential with a fourth order coefficient that becomes negative for $T > T_c$ (assuming that there are higher order terms that stabilize the potential).

In conclusion, there is no phase transition in this sector I.

- *Condensation in Sector II*

The Lagrangian is given by (2.46). It describes a complex scalar with $m^2 L^2 = -3$ and $qL = \sqrt{3}$. Comparing with the previous model, this scalar field has a greater mass and also greater charge. As discussed, increasing the mass lowers the critical temperature,

but increasing the charge raises it, so in this case it is not a priori obvious whether the critical temperature for this model will be lower or higher than the previous case. We have carried out the explicit calculation, including backreaction. We find that in this sector the critical temperature is $T_c \cong 0.70$, i.e. greater than $T_c \cong 0.56$ of sector I.

The results are shown in figure 6b, which reproduces the results of [6]. Figure 7 shows, in the same plot, the free energy for the three different solutions: the Reissner-Nordström AdS solution, the sector I hairy black hole representing retrograde condensation, and the sector II hairy black hole providing the dominant thermodynamics below $T_c \cong 0.70$.

In conclusion, among the two best candidates to condense in the three equal charge ensemble, the scalar mode of sector II is the one that condenses. Thus the consistent truncation carried out in [6] (particularized to the case when the Sasaki-Einstein manifold is S^5) seems to indeed pinpoint the relevant mode to study condensation in this ensemble. This is despite the fact that the scalar mode of sector II is not the one which is dual to the lowest dimension operator. The scalar mode of sector I is dual to a operator of lower dimension, but it gives rise to hairy black holes at $T > T_c$ with a free energy that is higher than the free energy of the RN black hole representing the uncondensed phase.

One important issue regards the stability of the $T = 0$ limit of the hairy black hole of sector II. In [20] it was pointed out that the the zero-temperature solutions are domain walls interpolating between non-supersymmetric AdS solutions at the horizon and AdS at infinity, which for compactifications based on spheres are unstable. This fact may be interpreted in different ways. On one hand, it might indicate that, in reaching $T = 0$, there must be a quantum phase transition to a stable fixed point of the $\mathcal{N} = 8$ potential. On the other hand, it could be that this sector just does not capture the dominant thermodynamic configuration even at finite temperature $T < T_c$, and that there is another sector within $\mathcal{N} = 8$ supergravity containing black hole solutions that at sufficiently low temperatures dominate the thermodynamics and smoothly flow to a supersymmetric fixed point (which would thus ensure stability).

In an attempt to search for such a sector with a stable $T = 0$ limit, we have examined sectors that include the $SU(2) \times U(1)$ supersymmetric fixed point of $\mathcal{N} = 8$ $SO(6)$ gauged supergravity to see if it is possible to have a flow towards this point. The simplest consistent truncation of this type is achieved in sector II with $A_1 = A_2$ and a different $A_3 \neq 0$, and setting $\tilde{\eta}_{2,3,4} = 0$ and $\varphi_2 = 0$ in the sector II Lagrangian (2.16), with the charges given by (2.36). The truncated Lagrangian is as follows

$$\mathcal{L} = \sqrt{-g} \left[R - \frac{1}{2}(\partial\varphi_1)^2 - \frac{1}{2}(\partial\eta)^2 - \frac{1}{2}\sinh^2(\eta)(\partial\theta + L^{-1}(2A_1 - A_3))^2 - \frac{1}{2}e^{-\frac{2\varphi_1}{\sqrt{6}}}F_{\mu\nu}^1 F^{1,\mu\nu} - \frac{1}{4}e^{\frac{4\varphi_1}{\sqrt{6}}}F_{\mu\nu}^3 F^{3,\mu\nu} - V \right], \tag{3.43}$$

with

$$V = \frac{e^{-\frac{4\varphi_1}{\sqrt{6}}}}{L^2} \cosh^2\left(\frac{1}{2}\eta\right) \left((1 + 2e^{\sqrt{6}\varphi_1}) \cosh(\eta) - 8e^{\frac{\sqrt{6}\varphi_1}{2}} - 1 - 6e^{\sqrt{6}\varphi_1} \right). \quad (3.44)$$

The resulting model indeed contains the $SU(2) \times U(1)$ supersymmetric fixed point as a classical solution of the equations and, moreover, one can write a consistent ansatz for black holes with three equal charges (despite the fact that $A_3 \neq A_1 = A_2$). The numerical resolution of this system presents new complications. To the extent we were able to carry out the numerical analysis reliably, we found no condensed phase in this sector, i.e. no black hole with three equal charges and $\tilde{\eta}_1$ hair with the required asymptotic behavior.

4 Condensation from other sectors

Since there are 42 physical scalars in the five-dimensional $\mathcal{N} = 8$ supergravity description, a natural question is whether there could be a scalar not considered in our analysis that could condense earlier. Clearly, an accurate answer to this point would require a long analysis. Nevertheless, one can guess which scalars could be relevant thermodynamically from the following analysis:

- From the $(\mathbf{1}, \mathbf{20}_S)$, there are 2 neutral scalars (namely φ_1 and φ_2) and the three complex scalars that we have considered in sector I with $U(1) \times U(1) \times U(1)$ charges $(\pm 2, 0, 0)$ and cyclic permutations. The remaining scalars are 6 complex scalars with charges $(\pm 1, \pm 1, 0)$ (signs unrelated) and cyclic permutations. These last scalars have the same $m^2 L^2 = -4$, but lower charges under each $U(1)$. However, they could compete with η_i in the two-equal charge thermal ensemble, since they would have the same charge under a diagonal $U(1)_D \subset U(1) \times U(1)$. In view of this possibility, we have explicitly investigated truncations including these charged scalar fields.

Starting from the symmetric scalar matrix T_{mn} , parametrising the $(\mathbf{1}, \mathbf{20}_S)$, one can perform truncations based on the following discrete $SO(6)$ transformations:

$$\left(\begin{array}{c} -\mathbb{I}_4 \\ +\mathbb{I}_2 \end{array} \right), \quad \left(\begin{array}{ccc} & 1 & \\ -1 & & \\ & -1 & \\ & & 1 \\ & & & 1 \\ & & & & -1 \end{array} \right), \quad \left(\begin{array}{ccc} & & 1 \\ & -1 & \\ -1 & & \\ & & & \mathbb{I}_2 \end{array} \right). \quad (4.1)$$

This leads to a parametrisation of T_{mn} that includes a dilaton φ_1 and a charged scalar field η (in the gauge where the corresponding angular variable θ is set to zero) of the

$Z_i = \eta_i e^{i\theta_i}$, $i = 1, 2, 3$, with mass $m^2 L^2 = -4$, have conformal dimension defined by the standard AdS/CFT relation $\Delta(\Delta - 4) = m^2 L^2 = -4$, i.e. $\Delta = 2$. In addition, they must carry charge = 2 under the relevant U(1). The natural candidate for the CFT operators which are dual to Z_i are

$$O_i = \text{Tr}[\Phi_i^2], \tag{5.1}$$

where Φ_i are the three chiral superfields of $\mathcal{N} = 4$ super Yang-Mills theory. These are BPS operators. The scalars of sector II, $\tilde{Z}_a = \tilde{\eta}_a e^{i\tilde{\theta}_a}$, are also dual to BPS operators whose relevant component is made with fermion bilinears [20]

$$\tilde{O}_a = \text{Tr}[\lambda_a \lambda_a] + \text{h.c.} . \tag{5.2}$$

Our analysis suggests that in the three-charge ensemble, the first operator to condense is \tilde{O}_4 , just as in the type IIB truncations of [6] (for the case of S^5). For the two charge ensemble, we have found no phase transition in the sectors we studied. Finally, in the one-charge case we have found an intriguing condensation (corresponding to a vacuum expectation value for O_3) where the hairy black hole rapidly loses its energy and reduces to a small size. We argued that this describes a new superconducting phase with an order parameter that becomes very large at some $T_1 < T_c$. The specific heat has first a jump at T_c — which is characteristic of second-order phase transitions — and then a very high peak near T_1 , where entropy and internal energy get suddenly very small. It would be interesting to identify condensed matter systems with similar properties.

Acknowledgments

We would like to thank Aristomenis Donos, Jerome Gauntlett, Maxim Mostovoy, Diego Rodriguez-Gomez and Paul Townsend for useful discussions. F.A. is supported by a MEC FPU Grant No. AP2008-04553. D.R. acknowledges support by a VIDI grant from the Netherlands Organisation for Scientific Research (NWO). J.R. acknowledges support by MCYT Research Grant No. FPA2007-66665 and Generalitat de Catalunya under project 2009SGR502.

Open Access. This article is distributed under the terms of the Creative Commons Attribution Noncommercial License which permits any noncommercial use, distribution, and reproduction in any medium, provided the original author(s) and source are credited.

References

- [1] S.S. Gubser, *Breaking an Abelian gauge symmetry near a black hole horizon*, *Phys. Rev. D* **78** (2008) 065034 [[arXiv:0801.2977](#)] [[SPIRES](#)].
- [2] S.A. Hartnoll, C.P. Herzog and G.T. Horowitz, *Holographic superconductors*, *JHEP* **12** (2008) 015 [[arXiv:0810.1563](#)] [[SPIRES](#)].
- [3] C.P. Herzog, *Lectures on holographic superfluidity and superconductivity*, *J. Phys. A* **42** (2009) 343001 [[arXiv:0904.1975](#)] [[SPIRES](#)].
- [4] G.T. Horowitz, *Introduction to holographic superconductors*, [arXiv:1002.1722](#) [[SPIRES](#)].

- [5] F. Denef and S.A. Hartnoll, *Landscape of superconducting membranes*, *Phys. Rev. D* **79** (2009) 126008 [[arXiv:0901.1160](#)] [[SPIRES](#)].
- [6] S.S. Gubser, C.P. Herzog, S.S. Pufu and T. Tesileanu, *Superconductors from superstrings*, *Phys. Rev. Lett.* **103** (2009) 141601 [[arXiv:0907.3510](#)] [[SPIRES](#)].
- [7] J.P. Gauntlett, J. Sonner and T. Wiseman, *Holographic superconductivity in M-theory*, *Phys. Rev. Lett.* **103** (2009) 151601 [[arXiv:0907.3796](#)] [[SPIRES](#)].
- [8] J.P. Gauntlett, J. Sonner and T. Wiseman, *Quantum criticality and holographic superconductors in M-theory*, *JHEP* **02** (2010) 060 [[arXiv:0912.0512](#)] [[SPIRES](#)].
- [9] A. Donos, J.P. Gauntlett, N. Kim and O. Varela, *Wrapped M5-branes, consistent truncations and AdS/CMT*, *JHEP* **12** (2010) 003 [[arXiv:1009.3805](#)] [[SPIRES](#)].
- [10] F. Aprile, D. Rodriguez-Gomez and J.G. Russo, *p-wave holographic superconductors and five-dimensional gauged supergravity*, *JHEP* **01** (2011) 056 [[arXiv:1011.2172](#)] [[SPIRES](#)].
- [11] P. Breitenlohner and D.Z. Freedman, *Positive energy in anti-de Sitter backgrounds and gauged extended supergravity*, *Phys. Lett. B* **115** (1982) 197 [[SPIRES](#)].
- [12] P.K. Townsend, *Positive energy and the scalar potential in higher dimensional (super)gravity theories*, *Phys. Lett. B* **148** (1984) 55 [[SPIRES](#)].
- [13] M. Günaydin, L.J. Romans and N.P. Warner, *Gauged N = 8 supergravity in five-dimensions*, *Phys. Lett. B* **154** (1985) 268 [[SPIRES](#)].
- [14] M. Pernici, K. Pilch and P. van Nieuwenhuizen, *Gauged N = 8 D = 5 supergravity*, *Nucl. Phys. B* **259** (1985) 460 [[SPIRES](#)].
- [15] A. Donos and J.P. Gauntlett, *Superfluid black branes in AdS₄ × S⁷*, [arXiv:1104.4478](#) [[SPIRES](#)].
- [16] L. Huijse and S. Sachdev, *Fermi surfaces and gauge-gravity duality*, [arXiv:1104.5022](#) [[SPIRES](#)].
- [17] E. Cremmer, *Supergravities in 5 dimensions*, invited paper at the *Nuffield Gravity Workshop*, Cambridge U.K., Jun 22–Jul 12, 1980 [[SPIRES](#)].
- [18] H.J. Kim, L.J. Romans and P. van Nieuwenhuizen, *The mass spectrum of chiral N = 2 D = 10 supergravity on S⁵*, *Phys. Rev. D* **32** (1985) 389 [[SPIRES](#)].
- [19] A. Khavaev and N.P. Warner, *A class of N = 1 supersymmetric RG flows from five-dimensional N = 8 supergravity*, *Phys. Lett. B* **495** (2000) 215 [[hep-th/0009159](#)] [[SPIRES](#)].
- [20] N. Bobev, A. Kundu, K. Pilch and N.P. Warner, *Supersymmetric charged clouds in AdS₅*, *JHEP* **03** (2011) 070 [[arXiv:1005.3552](#)] [[SPIRES](#)].
- [21] J.T. Liu, H. Lü, C.N. Pope and J.F. Vazquez-Poritz, *New supersymmetric solutions of N = 2, D = 5 gauged supergravity with hyperscalars*, *JHEP* **10** (2007) 093 [[arXiv:0705.2234](#)] [[SPIRES](#)].
- [22] M. Cvetič, S.S. Gubser, H. Lü and C.N. Pope, *Symmetric potentials of gauged supergravities in diverse dimensions and Coulomb branch of gauge theories*, *Phys. Rev. D* **62** (2000) 086003 [[hep-th/9909121](#)] [[SPIRES](#)].
- [23] E. Bergshoeff, M. Nielsen and D. Roest, *The domain walls of gauged maximal supergravities and their M-theory origin*, *JHEP* **07** (2004) 006 [[hep-th/0404100](#)] [[SPIRES](#)].

- [24] K. Behrndt, M. Cvetič and W.A. Sabra, *Non-extreme black holes of five dimensional $N = 2$ AdS supergravity*, *Nucl. Phys. B* **553** (1999) 317 [[hep-th/9810227](#)] [[SPIRES](#)].
- [25] G.T. Horowitz and M.M. Roberts, *Holographic superconductors with various condensates*, *Phys. Rev. D* **78** (2008) 126008 [[arXiv:0810.1077](#)] [[SPIRES](#)].
- [26] S.S. Gubser and F.D. Rocha, *Peculiar properties of a charged dilatonic black hole in AdS₅*, *Phys. Rev. D* **81** (2010) 046001 [[arXiv:0911.2898](#)] [[SPIRES](#)].
- [27] J. Erdmenger, V. Grass, P. Kerner and T.H. Ngo, *Holographic superfluidity in imbalanced mixtures*, [arXiv:1103.4145](#) [[SPIRES](#)].
- [28] D. Yamada and L.G. Yaffe, *Phase diagram of $N = 4$ super-Yang-Mills theory with R -symmetry chemical potentials*, *JHEP* **09** (2006) 027 [[hep-th/0602074](#)] [[SPIRES](#)].
- [29] T.J. Hollowood, S. Prem Kumar, A. Naqvi and P. Wild, *$N = 4$ SYM on S^3 with near critical chemical potentials*, *JHEP* **08** (2008) 046 [[arXiv:0803.2822](#)] [[SPIRES](#)].
- [30] A. Buchel and C. Pagnutti, *Exotic hairy black holes*, *Nucl. Phys. B* **824** (2010) 85 [[arXiv:0904.1716](#)] [[SPIRES](#)].
- [31] A. Buchel and A. Patrushev, *Can the correlated stability conjecture be saved?*, [arXiv:1102.5331](#) [[SPIRES](#)].
- [32] J.P. Kuenen, *Measurements on the surface of Van der Waals for mixtures of carbonic acid and methyl chloride*, *Commun. Phys. Lab. Univ. Leiden* **4** (1892).
- [33] D. Katz and F. Kurata, *Retrograde condensation*, *Ind. Eng. Chem.* **32** (1940) 817.
- [34] F. Aprile and J.G. Russo, *Models of holographic superconductivity*, *Phys. Rev. D* **81** (2010) 026009 [[arXiv:0912.0480](#)] [[SPIRES](#)].
- [35] S. Franco, A. Garcia-Garcia and D. Rodriguez-Gomez, *A general class of holographic superconductors*, *JHEP* **04** (2010) 092 [[arXiv:0906.1214](#)] [[SPIRES](#)].
- [36] F. Aprile, S. Franco, D. Rodriguez-Gomez and J.G. Russo, *Phenomenological models of holographic superconductors and Hall currents*, *JHEP* **05** (2010) 102 [[arXiv:1003.4487](#)] [[SPIRES](#)].
- [37] C.P. Herzog, *An analytic holographic superconductor*, *Phys. Rev. D* **81** (2010) 126009 [[arXiv:1003.3278](#)] [[SPIRES](#)].
- [38] S. Bhattacharyya, S. Minwalla and K. Papadodimas, *Small hairy black holes in $AdS_5 \times S^5$* , [arXiv:1005.1287](#) [[SPIRES](#)].